超新星爆発時における プレソーラーSiC粒子の形成

(Formation of presolar SiC grains in the ejecta of supernovae)

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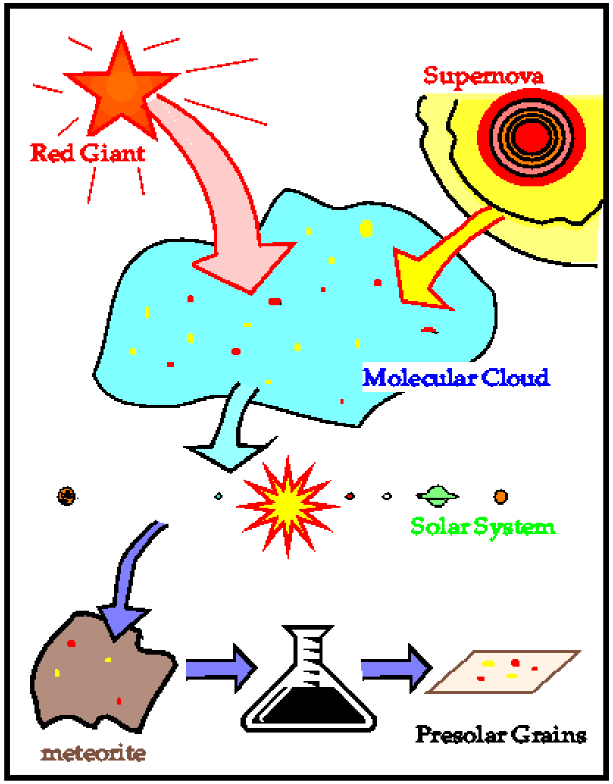
1-1. Presolar grains

○ Presolar grains

- discovered in primitive meteorites and IDPs
- showing peculiar isotopic compositions
  (different from the solar system's materials)
- thought to have formed in stars before the Sun was formed



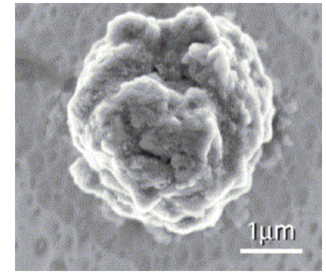
http://presolargrains.net/

➜ offering key information on

- nuclear processes in the parent stars
- physical conditions in which they formed

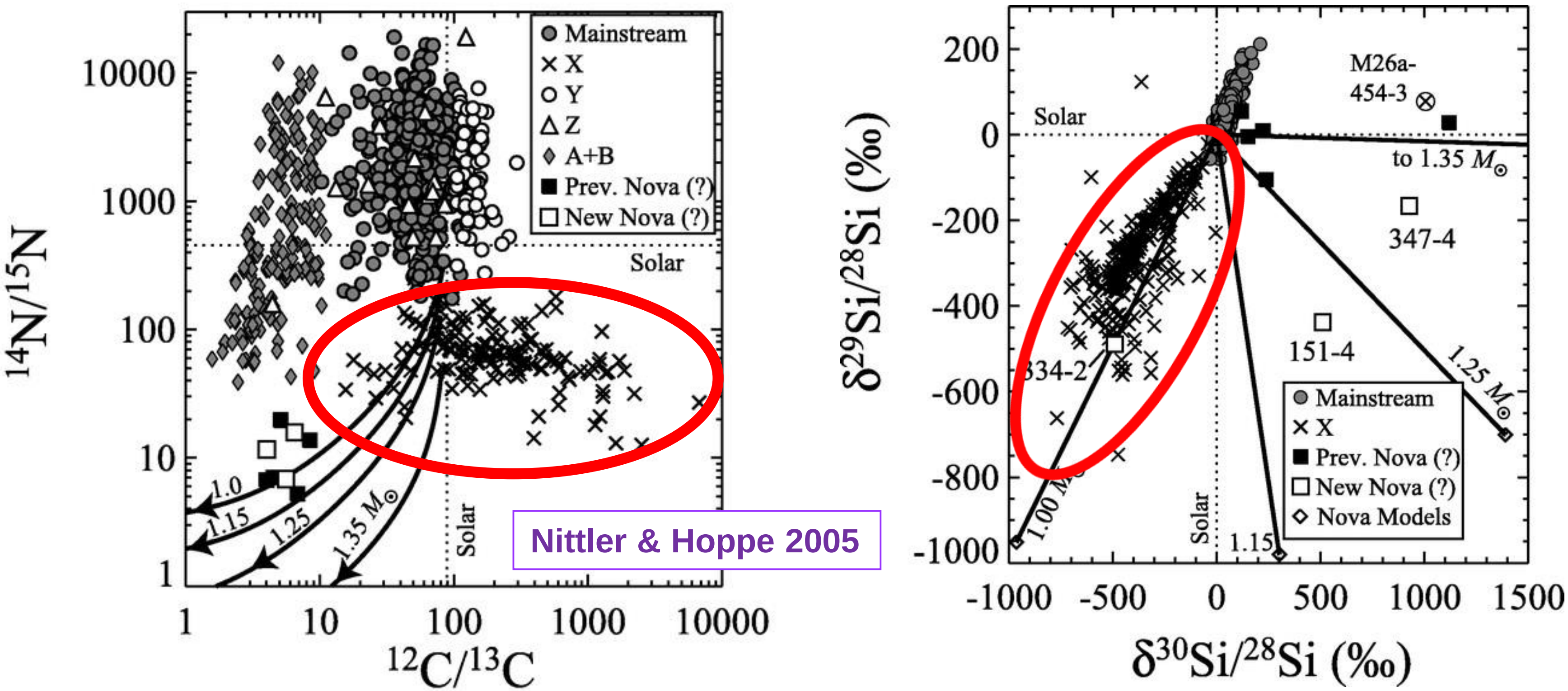
- abundance (volume fraction): ~100-500 ppm (~0.01-0.05 %)
- mineral composition
  graphite, SiC, TiC, Si_3N_4, Al_2O_3, $MgAl_2O_4$, Mg_2SiO_4, $MgSiO_3$ …

1-2. Presolar Type X SiC grains



Nittler 2003

○ Presolar SiC grains (~10 ppm)



Nittler & Hoppe 2005

○ Type X SiC grains (~0.1 ppm, size: 0.1-20 μm)

- 12C and 28Si rich, with excesses in 44Ca (44Ti), 49Ti (49V)

➜ originated from core-collapse supernovae

1-3. Long-lasting questions

○ Fact

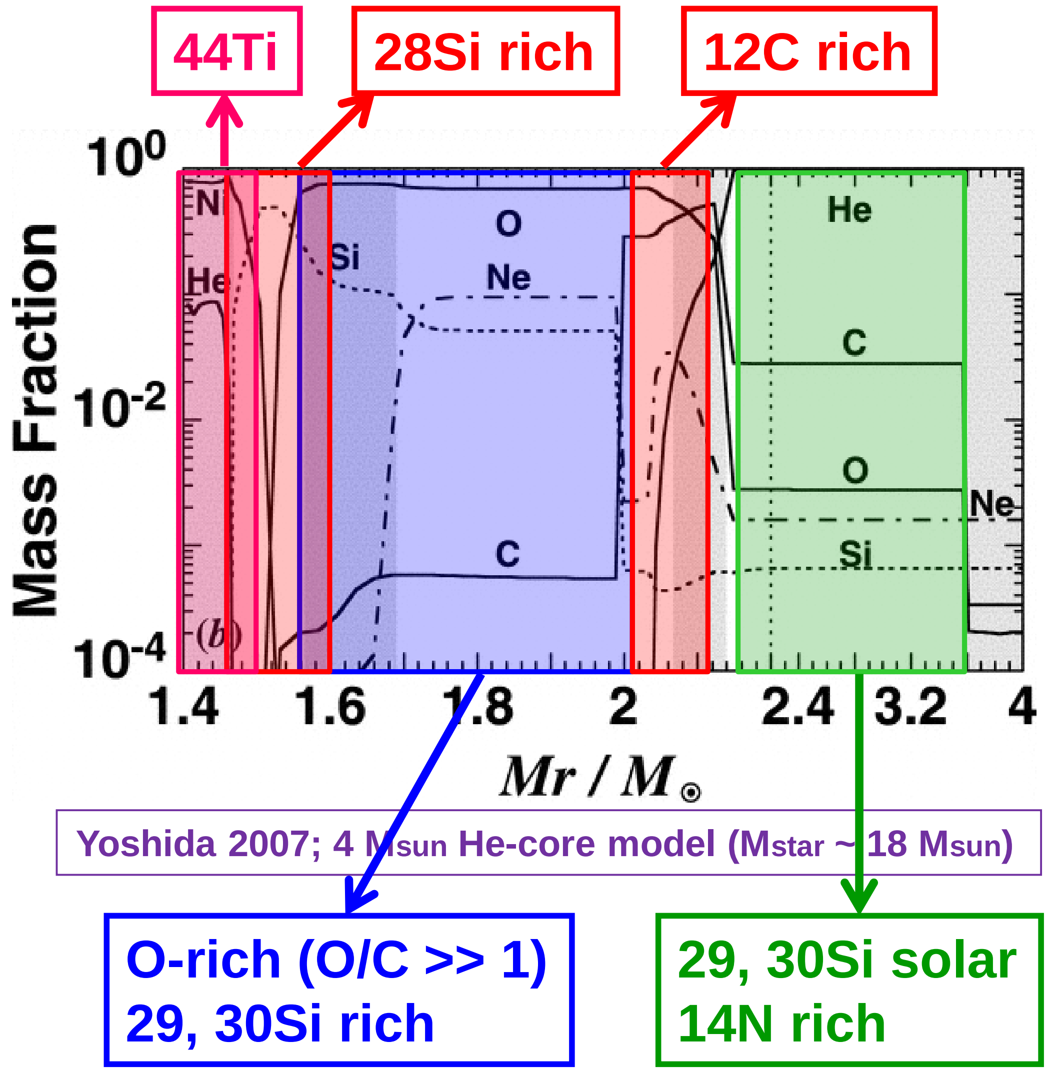
A fraction of presolar SiC grains is highly likely SN-origin

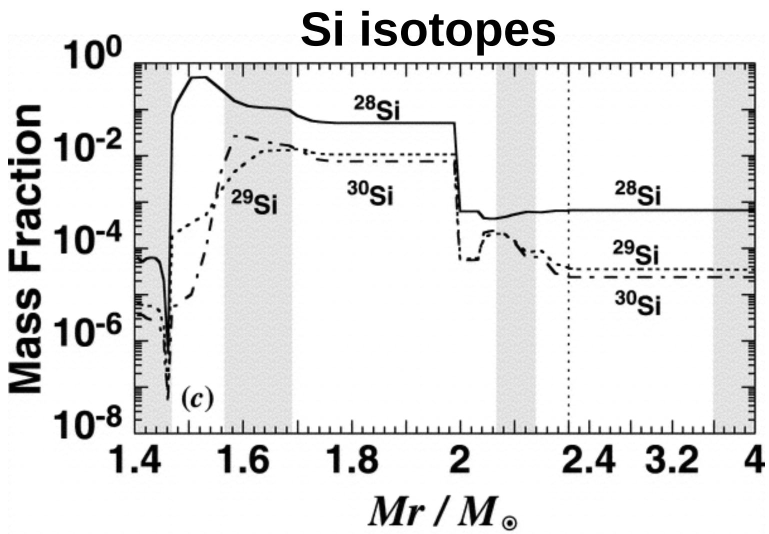
○ Question

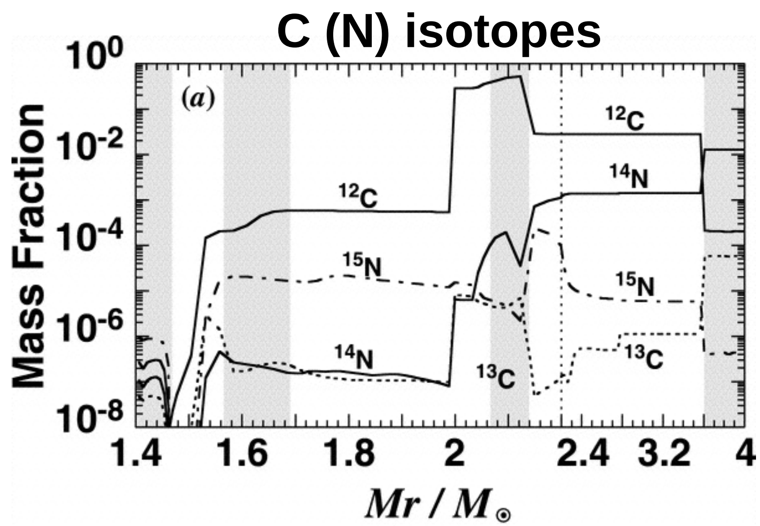
How do these SiC grains form in the ejecta of core-collapse supernovae?

Any theoretical studies have not yet realized the
formation of SiC grains in supernovae

2-1. Where SiC grains form in the ejecta?

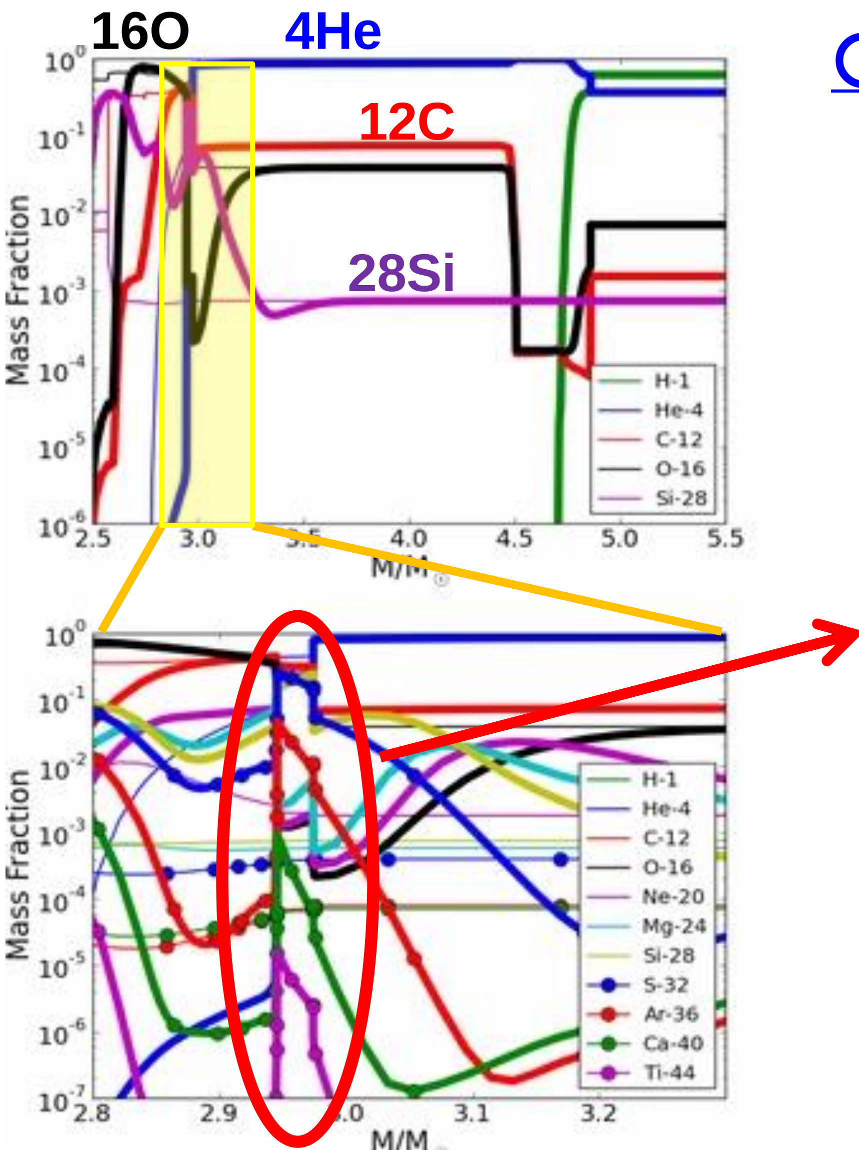






require special condition of ejecta mixing

2-2. Possible formation site of SiC grains in SNe



Pignatari+2013

○ Moderately energetic SNe

E_{kin} = (3-5)x10^{51} erg

(normal SNe : E_{kin} ~ 1x10^{51} erg)

explosive burning happens at the boundary between O-rich region and He-rich region

➜ producing 12C-rich, 28Si-rich, (44Ti-rich) region

3(4He) ➔ 12C

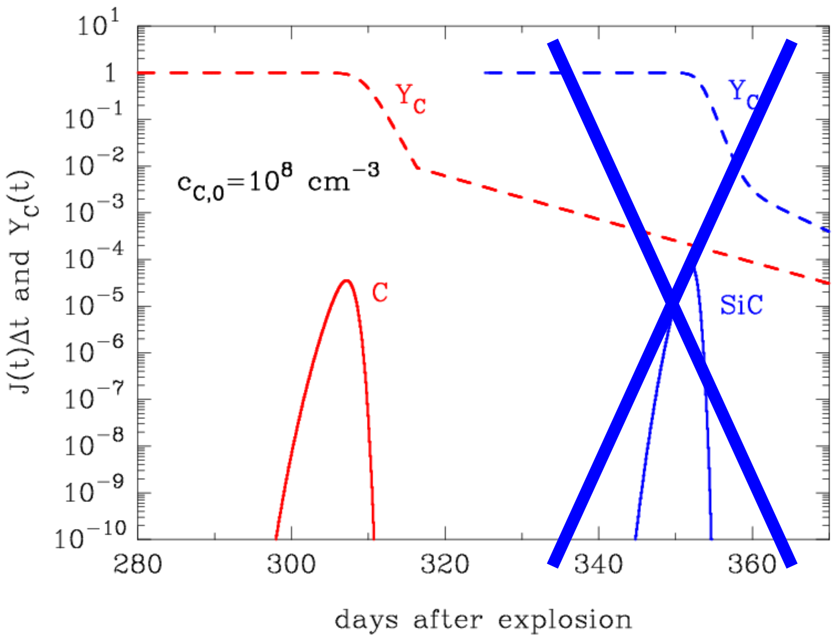
12C + n(4He) ➜ α-elements

2-3. No formation of SiC in the calculations

○ Condensation temperature

- C grains : Tcon = 1700-2200 K
- SiC grains: Tcon = 1300-1800 K

➜ C grains first condense to use up the gas-phase C atoms



○ Formation path of grains

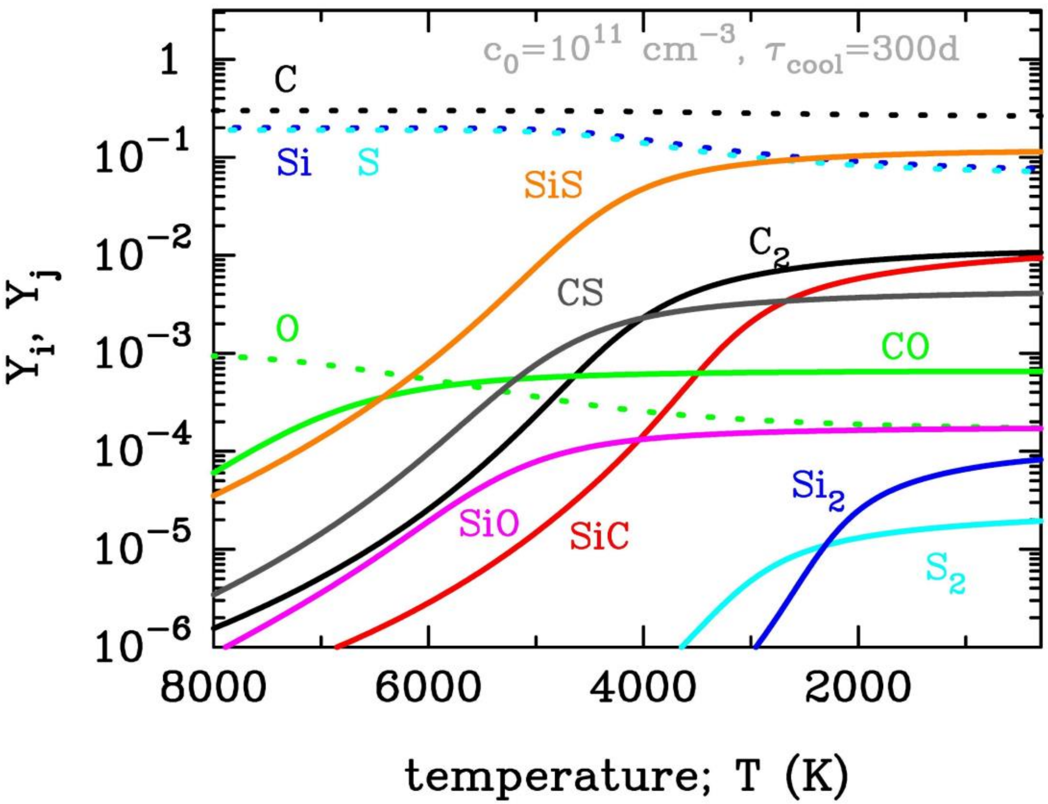
$$C_n + C \longrightarrow C_{n+1}$$

$$(SiC)_n + Si + C \longrightarrow (SiC)_{n+1}$$

$$(SiC)_n + SiC \longrightarrow (SiC)_{n+1}$$

This work

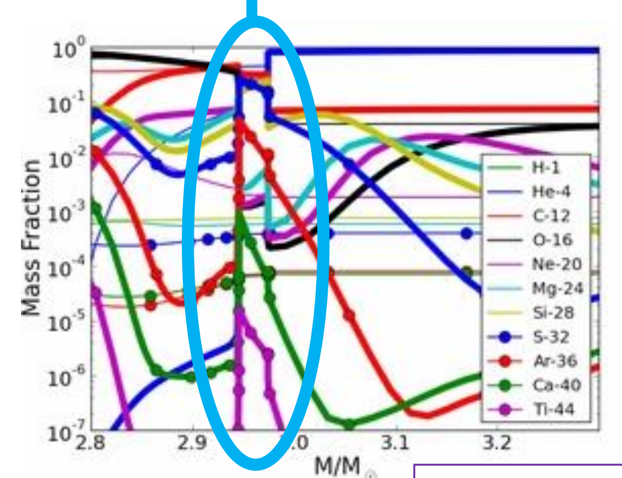
3-1. Formation of SiC molecules



Element	# abundance
He	0.3
C	0.3
O	0.001
Ne	0.001
Mg	0.002
Si	0.2
S	0.19
Ar	0.006

$$c(t) = c_{100}\left(\frac{t}{100\ \mathrm{day}}\right)^{-3}\ \mathrm{cm}^{-3}$$

$$T(t) = T_0 \exp\left(-\frac{t-t_0}{\tau_{\mathrm{cool}}}\right)$$



Pignatari+2013

3-2. Equation of dust formation

○ Master equations of dust growth

$$\frac{dc_n}{dt} = \frac{1}{2}\sum_{i=1}^{n-1} K_{i,n-i}(T)c_i\left(c_{n-i} - c_n e^{\gamma_{i,n-1}}\right) - c_n \sum_{l=1}^{n_{\max}} K_{n,l}(T)\left(c_l - c_{n+l}e^{\gamma_{l,n}}\right)$$

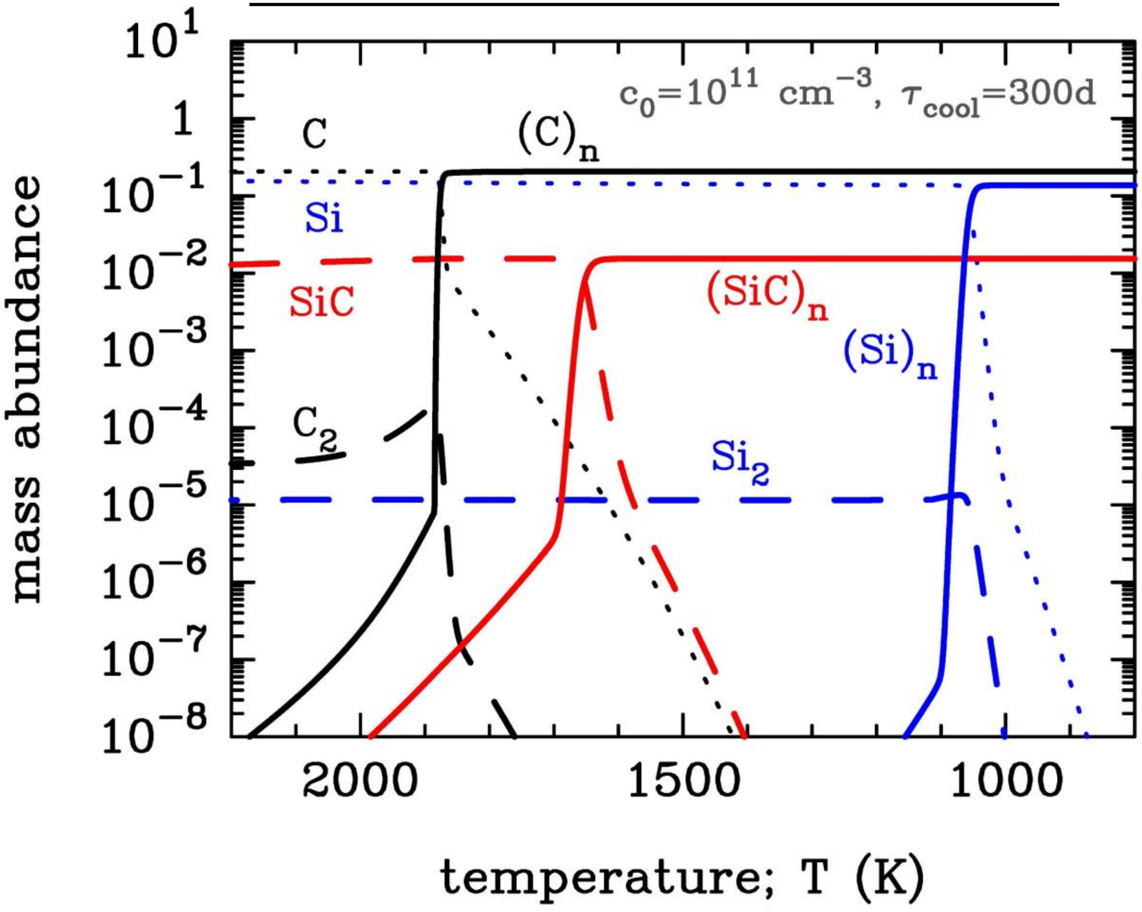
$$K_{i,j}(T) = \alpha_s \pi (a_i + a_j)^2 \left(\frac{8kT}{\pi m_{i,j}}\right)^{\frac{1}{2}}$$

$$\gamma_{i,n-i}(T) = \frac{4\pi a_0^2 \sigma}{kT}\left[(n-1)^{\frac{2}{3}} - (n-i-1)^{\frac{2}{3}} - (i-1)^{\frac{2}{3}}\right] - \ln S_i$$

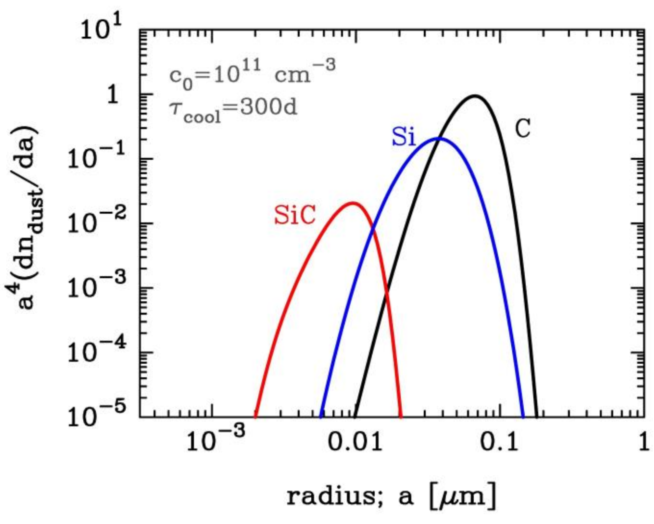
taking account of nucleation and grain growth through gas accretion and coagulation between grains simultaneously!

3-3. Can SiC grains form in the SN ejecta?

behavior of dust formation



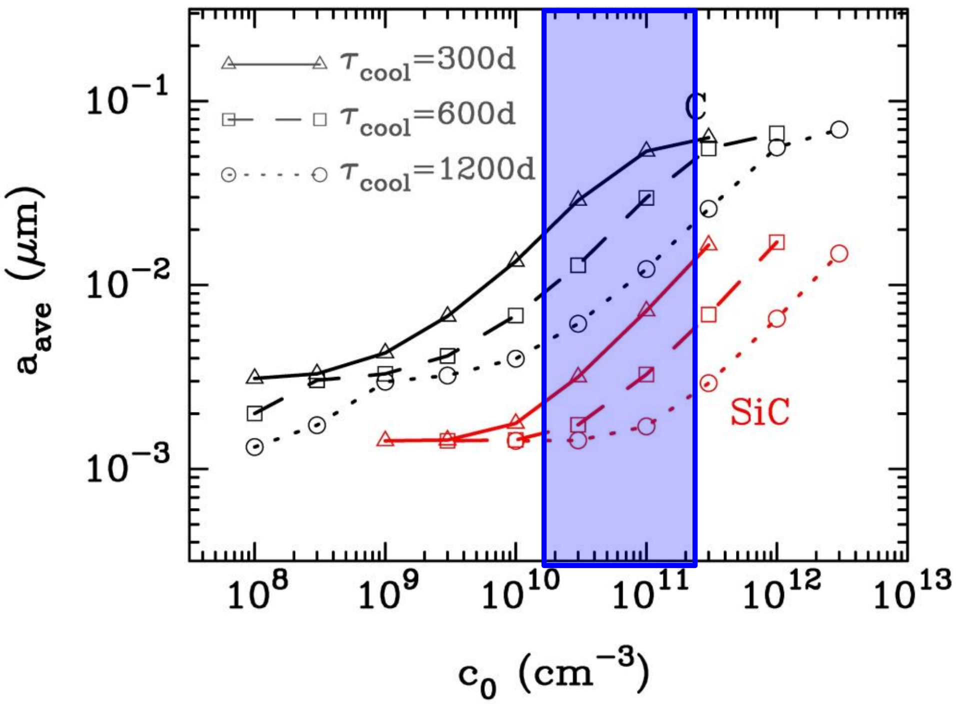
size distribution of dust



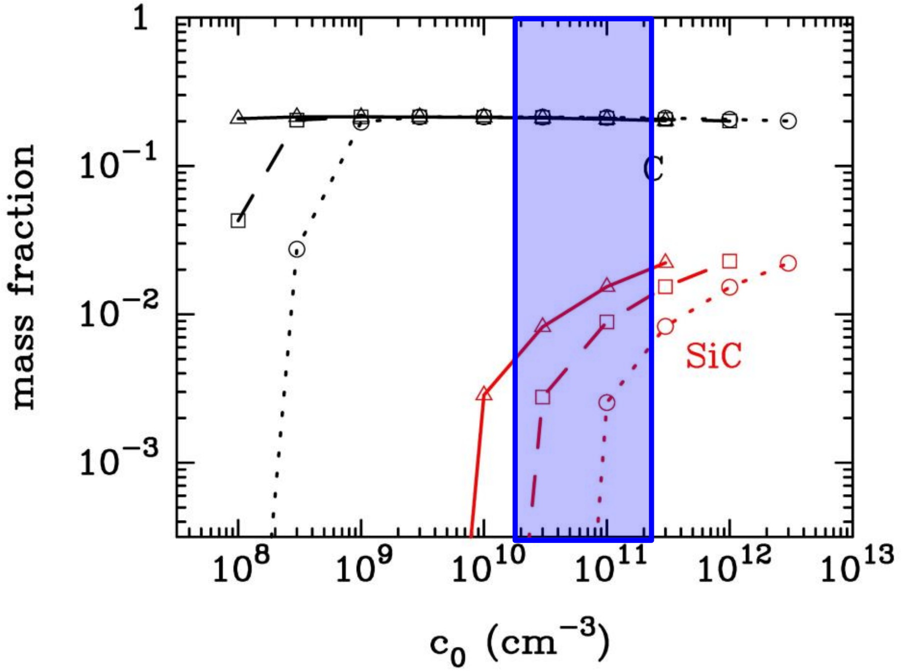
Yes, the formation of SiC grains is possible !

3-4. Dependence of grain radius on density

Average radius of dust grains



Mass fraction of dust grains



- **Radius of newly formed grains is larger for shorter τ_cool**
- **Radius and mass of SiC grains are lower than those of C grains by factors of ~5 and ~10, respectively.**

4. Summary of this talk

We investigate the formation of SiC grains in the ejecta of supernovae, self-consistently treating

- **formation of SiC molecules**
- **growth of SiC grains via accretion of molecules**
- **growth of SiC grains via grain-grain coagulation**

We have realized, for the first time, the formation of SiC grains in the ejecta of supernovae

➜ Formation of large presolar Type-X SiC grains above 0.1 μm is still difficult

Any idea, please !!